Constraining the epoch of reionization with variance statistic: simulations of the LOFAR case

Ajinkya H. Patil,[1][*] Saleem Zaroubi,[1] Emma Chapman,[2] Vibor Jelić,[1,3]
Filipe B. Abdalla,[2] Khan M. Asad,[1] Gianni Bernardi,[4] Michiel A. Brentjens,[3]
A. G. de Bruyn,[1,3] Sander Bus,[1] Benedetta Ciardi,[5] Soobash Daiboo,[1]
Elizabeth R. Fernandez,[1] Abhik Ghosh,[1] Geraint Harker,[6] Hannes Jensen,[7]
Sanaz Kazemi,[1] Léon V. E. Koopmans,[1] Panagiotis Labropoulos,[3] Maaijke Mevius,[1,3]
Oscar Martinez,[1] Garrelt Mellema,[7] Andre. R. Offringa,[1,8] Vishhambhar N. Pandey,[1,3]
Joop Schaye,[9] Rajat M. Thomas,[1] Harish K. Vedantham,[1] Vamsikrishna Veligatla,[1]
Stefan J. Wijnholds[3] and Sarod Yatawatta[1,3]

[1]*Kapteyn Astronomical Institute, University of Groningen, PO Box 800, 9700AV Groningen, the Netherlands*
[2]*Department of Physics & Astronomy, University College London, Gower Street, London WC1E 6BT*
[3]*ASTRON, PO Box 2, NL-7990AA Dwingeloo, the Netherlands*
[4]*SKA SA, 3rd Floor, The Park, Park Road, Pinelands, 7405, South Africa*
[5]*Max-Planck Institute for Astrophysics, Karl-Schwarzschild-Strasse 1, D-85748 Garching bei München, Germany*
[6]*Center for Astrophysics and Space Astronomy, University of Colorado, 389 UCB, Boulder, Colorado 80309-0389, USA*
[7]*Department of Astronomy and Oskar Klein Centre, Stockholm University, AlbaNova, SE-10691 Stockholm, Sweden*
[8]*Mount Stromlo Observatory, RSAA, Cotter Road, Weston Creek ACT 2611, Australia*
[9]*Leiden Observatory, Leiden University, PO Box 9513, NL-2300 RA Leiden, the Netherlands*

ABSTRACT

Several experiments are underway to detect the cosmic redshifted 21-cm signal from neutral hydrogen from the Epoch of Reionization (EoR). Due to their very low signal-to-noise ratio, these observations aim for a statistical detection of the signal by measuring its power spectrum. We investigate the extraction of the variance of the signal as a first step towards detecting and constraining the global history of the EoR. Signal variance is the integral of the signal's power spectrum, and it is expected to be measured with a high significance. We demonstrate this through results from a realistic, end-to-end simulation and parameter estimation pipeline developed for the Low Frequency Array (LOFAR)-EoR experiment. We find the variance statistic to be a promising tool and use it to demonstrate that the LOFAR can detect the EoR in 600 h of integration. Additionally, the redshift (z_r) and duration (Δz) of reionization can be constrained assuming a parametrization. We used an EoR simulation of $z_r = 7.68$ and $\Delta z = 0.43$ to test the pipeline, and were able to extract the parameters as $z_r = 7.72^{+0.37}_{-0.18}$ and $\Delta z = 0.53^{+0.12}_{-0.23}$ with 600 h of integration.

Key words: dark ages, reionization, first stars – methods: statistical – techniques: interferometric

1 INTRODUCTION

Advances in observational cosmology over the past century have made it possible to look very far out into the Universe. However, there still remains a big observational gap between the Cosmic Microwave Background (CMB) ($z \approx 1100$) and the low-redshift Universe ($z < 6$). An important global transition is expected to have occurred towards the end of this era, called the Epoch of Reionization (EoR). The first stars and galaxies formed during this epoch, and hydrogen in the Universe was reionized from their radiation after having being neutral for about 400 Myr.

Unfortunately, a dearth of observations makes the EoR a poorly constrained epoch. Current constraints are based on indirect observations of the high-redshift inter-galactic

[*] E-mail: patil@astro.rug.nl

medium (IGM), namely, quasar spectra (Fan et al. 2003, 2006), CMB polarization anisotropy (e.g. Hinshaw et al. 2013), the kinetic Sunyaev-Zel'dovich effect (Zahn et al. 2012), IGM temperature measurements (Theuns et al. 2002; Bolton et al. 2010), high-redshift galaxy surveys (e.g. Finkelstein et al. 2012), high redshift gamma ray bursts (Wang 2013), Lyman break galaxies (Pentericci et al. 2011; Ono et al. 2012; Schenker et al. 2012). However, redshifted 21-cm emission from neutral hydrogen has the potential to directly probe the IGM and hence the process of reionization. Therefore, many ongoing experiments aim to observe the EoR with low-frequency radio telescopes such as the Low Frequency Array (LOFAR) (van Haarlem et al. 2013), the Murchison Widefield Array (MWA) (Tingay et al. 2013), the Precision Array to Probe the Epoch of Reionization (PAPER) (Parsons et al. 2010), the Giant Meterwave Radio Telescope (GMRT) (Pen et al. 2008).

Detection of the EoR signal i.e. the redshifted 21-cm signal from the era of reionization is very challenging even with the new generation of radio telescopes. This is because in the redshift range of 6 to 10, the expected signal is only about $\sim$10 mK (at a resolution of 3 arcmin), whereas, the Galactic and extragalactic foregrounds are about $\sim$1 K (Bernardi et al. 2009, 2010). Moreover, even if the foregrounds would be perfectly removed, after hundreds of hours of integration, the system noise would still be an order of magnitude larger than the signal. Therefore, the current experiments aim for a statistical detection of the EoR rather than mapping the neutral hydrogen. This requires development of the statistical techniques to estimate the reionization parameters from noisy data. A commonly studied technique is power spectrum analysis (e.g. Morales & Hewitt 2004; McQuinn et al. 2006; Bowman, Morales & Hewitt 2006; Harker et al. 2010; Beardsley et al. 2013). Another possible statistic is the signal variance, which is the integral of signal's power spectrum. The variance (or root mean square) statistics of the EoR signal has been studied theoretically by Iliev et al. (2008); Jelić et al. (2008); Thomas et al. (2009); Harker et al. (2009); Bittner & Loeb (2011). Watkinson & Pritchard (2013) have shown that the current EoR experiments will struggle to distinguish the reionization models using the power spectrum, but are capable of breaking this degeneracy with the variance. In this paper, we investigate the measurement of signal variance with the LOFAR considering various instrumental parameters. And we use the variance as a quantitative measure to constrain the global history of reionization in early stage EoR experiments.

We have developed an end-to-end simulation pipeline to test the variance statistic in the case of the LOFAR-EoR experiment. The pipeline generates realistic mock observations by simulating the cosmic signal, foregrounds and noise. Efforts have been made to model as many instrumental effects as possible, except for calibration errors. For instance, we have considered the issue of primary beam effects on foreground removal and subsequent analysis. The pipeline also incorporates measurement of the EoR parameters, namely, the redshift and duration of the EoR.

The paper is organized as follows: In Section 2, we discuss our parametrization of the variance of the EoR signal as a function of redshift. In Section 3, we describe the simulation pipeline we have developed. The measurement of the signal variance and parameter estimation is discussed in Section 4. We show the results and demonstrate that LOFAR can constrain the EoR in 600 h in - Section 5, before summarizing our conclusions in Section 6.

2   PARAMETRIZATION

The observable quantity of the redshifted 21-cm emission is the differential brightness temperature δT_b i.e. the contrast between the 21-cm brightness temperature and the background CMB temperature T_CMB. At a given position in sky, δT_b is given by (Field 1959; Madau, Meiksin & Rees 1997; Furlanetto, Oh & Briggs 2006)

$$\delta T_\mathrm{b} \approx 9\, x_\mathrm{HI}(1+\delta)(1+z)^{\frac{1}{2}}\left[1 - \frac{T_\mathrm{CMB}(z)}{T_\mathrm{S}}\right]$$
$$\times \left[\frac{H(z)/(1+z)}{dv_\parallel/dr_\parallel}\right]\mathrm{mK}, \tag{1}$$

where δ is the cosmological mass density contrast, x_HI is the neutral hydrogen fraction, T_S is the spin temperature, $H(z)$ is Hubble parameter and $dv_\parallel/dr_\parallel$ is gradient of the proper velocity along the line of sight. Whenever we mention the EoR signal, we refer to the differential brightness temperature of the 21-cm radiation from reionization.

An interferometer can measure spatial fluctuations of δT_b as a function of frequency, or equivalently of cosmic redshift. The spatial fluctuations at a given redshift can be characterized by the power spectrum $P[\boldsymbol{k}]$ as

$$P[\boldsymbol{k}] = \delta T_b[\boldsymbol{k}]\delta T_b^*[\boldsymbol{k}] \tag{2}$$

where δT_b is measured at discrete values of wavenumber $\boldsymbol{k}$. The variance of the signal is given by

$$\mathrm{Var}(\delta T_b) = \langle P[\boldsymbol{k}]\rangle. \tag{3}$$

Our interest here lies in measuring the variance of the signal and its evolution with redshift.

Fig. 1 shows the evolution of the signal variance as predicted by the simulation code 21CMFAST (Mesinger, Furlanetto & Cen 2011). At the highest redshifts, the Universe is mostly neutral ($x_\mathrm{HI} \approx 1$), hence δT_b is driven by the cosmological density fluctuations δ. The density fluctuations grow with time to form the first ionizing sources, which then start to reionize their surrounding regions. This patchy nature of reionization leads to a rise in the variance of δT_b. The variance reaches its peak when approximately half of the Universe is ionized, but decreases thereafter. Eventually, it reaches zero as the entire Universe is reionized. The different curves in Fig. 1 are for different spatial resolutions and show that the observed variance depends on the resolution, or equivalently on the range of wavenumbers measured in the observed volume. Also, the higher the resolution, the earlier the variance peaks (Iliev et al. 2008). This is because higher resolution data are sensitive to smaller scale structures.

In order to learn about the process of reionization from variance measurements, we need a parametric model which describes the variance of δT_b in terms of the EoR model parameters. In this paper, we assume a model which enables us to constrain two important EoR parameters: the redshift of reionization z_r, defined as the redshift at which the variance of δT_b is maximum, and the duration of reionization Δz. The model is given by
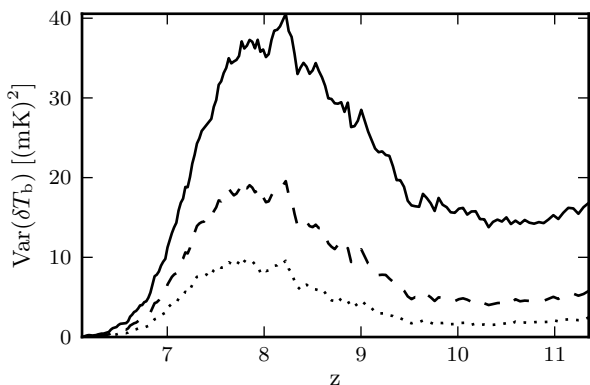
Figure 1. Evolution of variance of the epoch of reionization signal i.e. the 21-cm differential brightness temperature ($\delta T_{\rm b}$) with redshift, as predicted by simulation 21CMFAST. From top to bottom the curves correspond to decreasing spatial resolutions of 3 (resolution of the LOFAR core i.e. the central 2 km of the array at 150 MHz, $z \approx 8.5$), 7 and 12 arcmin. The resolution is expressed as the full width half maximum (FWHM) of the Gaussian kernel that was used to smooth the simulation maps.

$$\mathrm{Var}(\delta T_{\rm b}) = A\, f(z) \left(\frac{z}{z_0}\right)^{\beta}, \qquad (4)$$

where A is the scaling amplitude, $\beta\ (< 0)$ is the index of the (decaying) power law that the variance asymptotes to at high redshift and z_0 is the redshift which defines the regime $z \gg z_0$ in which the power law becomes dominant. The model is inspired by the fact that at high redshift, $\delta T_{\rm b}$ is driven by $(1 + \delta)$, which linear perturbation theory predicts to evolve as a power law. The function $f(z)$ describes the low-redshift behaviour of the signal and is defined as

$$f(z) = \frac{1}{2} \left[1 + \tanh\left(\frac{z - z_0}{\Delta z}\right)\right]. \qquad (5)$$

The redshift of reionization z_r is the redshift at which the variance reaches its maximum. Therefore, it is computed using the condition

$$\left.\frac{d\mathrm{Var}(\delta T_{\rm b})}{dz}\right|_{z_r} = 0. \qquad (6)$$

We translate the parameter z_0 to z_r by computing the difference z_{bias} between the two and then correcting for it as

$$z_r = z_0 + z_{bias}. \qquad (7)$$

Fig. 2 shows a model fit to the signal variance from a simulation. It can be seen that the model describes the simulation results well, except for the dip in the variance at high redshifts ($z \approx 10$). Such a decrease in the variance is expected to occur at the beginning of reionization (Iliev et al. 2012). The first objects form in density peaks and reionize their surrounding regions, which appear as holes in neutral hydrogen maps. These holes reduce the signal variance contributed by the corresponding density peaks. However, when many ionizing objects start to form, the variance is driven by the distribution of $x_{\rm HI}$ rather than by the density fluctuations. Therefore, the variance increases after the initial dip.

The small decrease in the variance at the beginning of reionization is not described by our parametrization. However, the LOFAR system noise increases rapidly at lower fre-
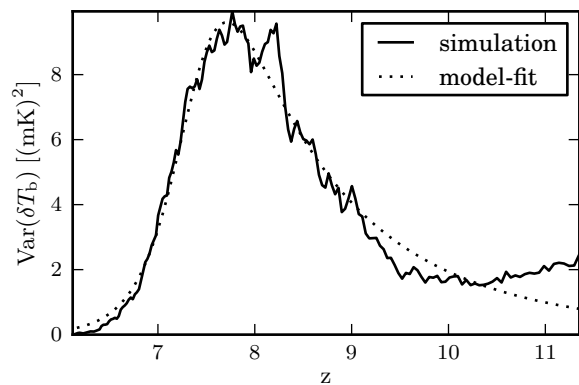

Figure 2. Fit of the model described by equations (4) and (5) to the signal variance predicted by the simulation at 12 arcmin resolution. The model describes the curve well except at $z > 9$.
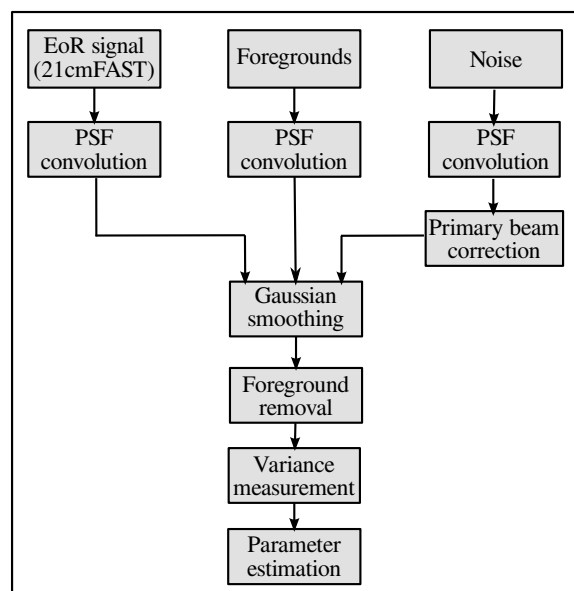

Figure 3. Block diagram of the end-to-end simulation and analysis pipeline.

quencies (higher redshifts), making the fitting performance by models which would incorporate this feature indistinguishable.

3 SIMULATION AND SIGNAL EXTRACTION PIPELINE

We generated mock observational data sets by adding simulations of the cosmological signal, foregrounds and noise. A data cube consisted of 170 frequency maps between 115 and 199.5 MHz (i.e. $z = 6$ to 11.4) at an interval of 0.5 MHz. Each frequency map initially represented a $10^{\circ} \times 10^{\circ}$ window with 1.17 arcmin resolution but was later corrected for the LOFAR field of view as will be discussed in Section 3.3. The important blocks of the simulation and signal extraction pipeline are described in the following subsections (please see Fig. 3 for a block diagram of the pipeline.)

3.1 The signal

Cosmological simulations of size larger than 600 comoving Mpc are required to simulate the field of view of LO-FAR. Full radiative transfer simulations on such large scales are computationally expensive. Instead, we used the semi-analytic code 21CMFAST (Mesinger, Furlanetto & Cen 2011; Mesinger & Furlanetto 2007) to simulate the EoR signal. 21CMFAST treats physical processes with approximate methods, but on scales larger than 1 Mpc its results are in good agreement with hydrodynamical simulations (Mesinger, Furlanetto & Cen 2011). The cosmological simulation used here is the same as in Chapman et al. (2012). The simulation was initialized with 1800^3 dark matter particles at $z = 300$. The code evolves the initial density and velocity fields to the redshifts of the EoR using linear perturbation theory. The velocity field used to perturb the initial conditions and the evolved simulation boxes were formed on a coarser grid of 450^3 and then interpolated up to 512^3. 21CMFAST uses the excursion set formalism to form dark matter haloes. We define the threshold for haloes contributing ionizing photons to be $10^9 M_\odot$. Once the evolved density, velocity and ionization fields have been obtained, the code computes the δT_b box at each redshift based on equation (1). Redshift space distortions were taken into account in our run, but we neglected spin temperature fluctuations by assuming $T_S \gg T_{CMB}$, i.e. the neutral gas has been heated well above the CMB for redshifts 6 to 12 (Pritchard & Loeb 2008). We combined the δT_b boxes at different redshifts using the method described by Thomas et al. (2009) to form an observational cube. An observational cube represents the 2D position on the sky and the third dimension corresponds to observation frequency or redshift.

3.2 Foregrounds

We used the simulations by Jelić et al. (2008, 2010) to model the foreground contamination. These simulations consider the following contributions:

(i) Galactic diffuse synchrotron emission (GDSE) due to the interaction of cosmic ray electrons with the galactic magnetic field. The GDSE is modelled as a power law as a function of frequency with the spectral index of -2.55 ± 0.1 (Shaver et al. 1999). The intensity and the spectral power law index of the GDSE are spatially modelled as Gaussian random fields. The power spectrum of these fields is assumed to be a power law with 2D index of -2.7. The mean brightness temperature at 120 MHz is 253 K with the standard deviation of 1.3 K.
(ii) Galactic localized synchrotron emission from supernova remnants (SNRs). Eight SNRs are placed randomly in the $10° \times 10°$ observational window. In order to model the extended nature of SNRs, they are modelled to be extended discs. Their angular size, flux density and spectral index are randomly chosen from the Green (2006) catalogue of the observed radio SNRs.
The combined galactic diffuse and localized synchrotron emission is the dominant component ($\sim$70 per cent) of the foregrounds at 100-200 MHz.
(iii) Galactic diffuse free-free emission due to bremsstrahlung radiation from diffuse ionized gas. It is modelled in a similar manner as the GDSE but the

frequency spectral index is fixed to -2.15 across the map. It contributes $\sim$1 per cent of the total foreground emission.
(iv) Unresolved extragalactic sources such as radio galaxies and clusters, contributing $\sim$27 per cent of the foreground emission. The simulated radio galaxies have power law spectra and random walk based clustering. The radio clusters have spectral indices of about -3 and are based on the cluster catalogue from the Virgo Consortium[1].

We assumed that calibration removes the bright point sources and hence these were not included in the foreground simulations.

3.3 Instrumental response and noise

Unlike the EoR and foreground simulations, an interferometer does not directly map the surface brightness distribution in the sky. Instead, it measures correlations of electric fields between pairs of interferometric elements (LO-FAR stations). These correlations are called visibilities. Each visibility $V(u, v)$ probes a certain spatial scale of the sky brightness distribution corresponding to the baseline (u, v) between a pair of stations. Therefore, the brightness distribution $I(l, m)$ on the sky can be mapped by taking the Fourier transform of the visibilities (Taylor, Carilli & Perley 1999) as given by

$$I_\nu(l, m)A_\nu(l, m) = \int V_\nu(u, v)\, e^{i2\pi(ul+vm)} du\, dv, \qquad (8)$$

where $A_\nu(l, m)$ is the primary beam response of the telescope, l and m are the direction cosines and the subscript ν indicates the frequency of the measurement. But in reality, an interferometer can sample only certain discrete values (u_k, v_k) of the continuous uv plane. The measured set is called as the sampling function $S_\nu(u, v)$ and is defined as

$$S_\nu(u, v) = \sum_k \delta^D(u - u_k, v - v_k), \qquad (9)$$

where δ^D is the 2-dimensional Dirac delta function. Additionally, each measured visibility contains a noise component $N_\nu(u, v)$. Therefore, the actual obtained image can be described as

$$I_\nu(l, m)A_\nu(l, m) + n_\nu(l, m) = \\ \int S_\nu(u, v)V'_\nu(u, v)\, e^{i2\pi(ul+vm)} du\, dv, \qquad (10)$$

where, $n_\nu(l, m)$ is the noise in the image space and $V'_\nu(u, v) = V_\nu(u, v) + N_\nu(u, v)$. We use uniform weighting, i.e. if a point in the uv plane is sampled multiple times in the sampling function, all the samples at that point are averaged. Therefore, the root-mean-square (RMS) noise in the gridded uv plane is inversely proportional to $\sqrt{S(u, v)}$. In this manner, the sampling function defines the power spectrum of the noise in an interferometric measurement.

In order to obtain realistic simulations of the noise, we filled the real and imaginary parts of the visibilities $N_\nu(u, v)$ with Gaussian random numbers. Visibilities were then Fourier transformed to the image space to obtain noise maps $n_\nu(l, m)$. By simulating the noise in this manner, we
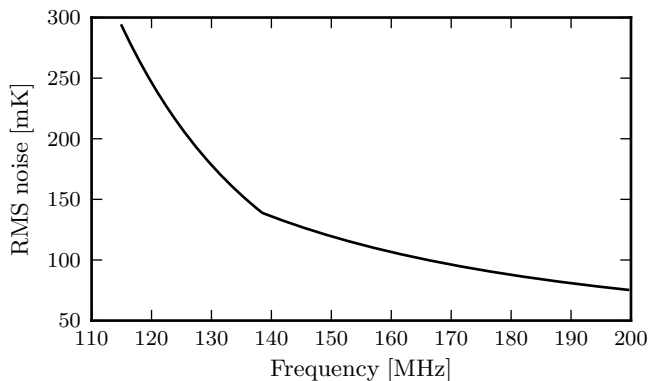
---

[1] http://www.mpa-garching.mpg.de/galform/virgo/hubble/

Figure 4. Root Mean Square (RMS) system noise as a function of frequency after 600 h and 0.5 MHz integration, used for normalizing the noise simulations.

incorporated the realistic power spectrum of the noise in our simulations. The noise maps were normalized to have the appropriate RMS values. Based on the theoretical calculations of the system equivalent flux density (Labropoulos et al. 2009), we expect the RMS noise to be about 120 mK at the resolution of 3 arcmin (i.e. the full resolution offered by the LOFAR core), at 150 MHz, after 600 h and 0.5 MHz of integration for uniformly weighted data. Fig. 4 shows the RMS noise used for normalizing the simulations as a function of frequency.

3.4 uv-mask or PSF convolution

An interferometer can only sample the spatial scales corresponding to its uv coverage. The effect of the uv coverage is equivalent to convolution with the point spread function (PSF) in the image plane. In order to mimic the effect of the PSF convolution, the simulated EoR, foreground and noise maps were Fourier transformed to the uv plane, multiplied by the uv coverage and Fourier transformed back to the image plane.

The uv coverage of an interferometer is frequency dependent because u and v are expressed in wavelengths. If a uv point is only sampled in a part of the bandwidth, it could introduce discontinuities in the foregrounds and noise properties along the frequency dimension. This would affect the performance of the foreground removal algorithms which are based on the assumption that the foregrounds are spectrally smooth. In order to avoid such discontinuities, we maintained the same uv coverage throughout the bandwidth. We achieved this by masking the intersections of the uv coverages at all frequencies. In other words, the uv points that were only partially covered in the bandwidth were discarded. Our uv mask allowed baselines between 40 and 800 wavelengths. We only considered stations within the central core of LOFAR because these densely sample the corresponding part of the uv plane. Stations outside the core provide longer baselines and are used in actual observations to remove the point sources during the calibration. They would then be discarded in the subsequent analysis.

The noise maps were simulated in the uv plane and already contain the uv coverage. However, they were also

multiplied by the uv mask to maintain the same uv coverage at all frequencies.

3.5 Primary beam correction

The image formed by an interferometer is the sky brightness distribution multiplied with its primary beam response as described by equation (8). Due to the primary beam response of the telescope, the strength of any observed signal from the sky (EoR and foregrounds) decreases away from the pointing direction. But noise, being uncorrelated among the visibilities, remains unaffected by the primary beam response. Hence the signal-to-noise (SNR) decreases away from the direction of pointing.

The primary beam response scales with wavelength. In the case of LOFAR, the Full Width Half Maximum (FWHM) of the primary beam changes from 4.75 degrees at 120 MHz to 2.85° at 200 MHz (van Haarlem et al. 2013). We find that the performance of the foreground removal suffers severely due to this frequency dependence, as shown further in Section 3.7. The primary beam correction reconstructs the frequency coherence of the foregrounds and hence improves the foreground removal. As a result of the correction, the EoR signal and the foregrounds have the same strength throughout the image but the noise increases towards the edges. We multiplied the noise maps by the reciprocal of the primary beam response $1/A_\nu(l, m)$ to include the beam correction in our simulations. We assumed a Gaussian primary beam with the same FWHM as that of the measured response.

In reality, the primary beam response resembles the $sinc^2$ function and its correction requires division by zero around the nulls. However, a Gaussian is a good approximation of the primary beam within the first null and we restrict the image size to be well within the first null. One way to avoid the primary beam correction in the case of actual observations and still get desirable foreground removal, could be to maintain the same primary beam shape throughout the bandwidth. This could be achieved by convolving of the visibilities with an appropriate kernel. However, it would restrict the field of view to the smallest possible case i.e. that obtained at the highest observation frequency. A better alternative would be to incorporate the beam model in the foreground removal algorithm. Current efforts are focused on this front and we leave this topic for a future paper. We would like to note that the treatment of the primary beam in this paper is very simple, but in reality, the beam response can change with time and differ for different LOFAR stations. More detailed models of the primary beam may be required for the analysis of the actual observed data.

3.6 Gaussian smoothing

The expected EoR signal RMS at the full resolution offered by the LOFAR core (~3 arcmin) is about 6 mK at 150 MHz, whereas the noise RMS is 120 mK after 600 h, 0.5 MHz integration. Therefore, the SNR at 3 arcmin resolution is very low. Not only may the signal detection be extremely difficult with such poor SNR, but the foreground removal may also be ineffective with such noisy data. The reason for poor SNR is the higher noise contribution at small spatial scales, which correspond to few long baselines. Even within the LOFAR
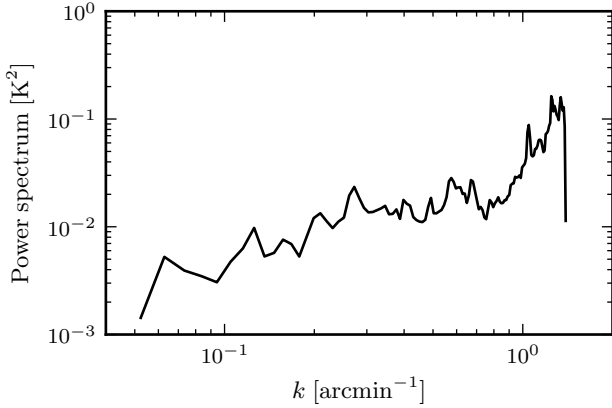
Figure 5. Radially averaged 2-dimensional power spectrum of the noise at 150 MHz as a function of wavenumber $k = 2\pi/\theta$, where θ is the angular scale. Most of the noise power is concentrated at large k values due to the lower sampling density of the outer part of uv plane.
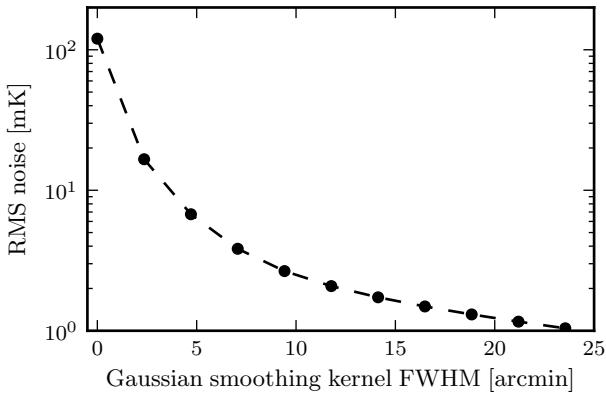

Figure 6. RMS noise as a function of smoothing scale at 150 MHz after 600 h and 0.5 MHz integration. The smoothing scale of zero means no smoothing. The noise initially decreases rapidly with smoothing as we effectively down-weight the long baselines.

core, the longer baselines are fewer in number, causing lower sampling density in the outer part of the uv coverage. Therefore the noise power is mostly concentrated on small spatial scales, as shown in Fig.5. We took advantage of this fact to reduce the noise significantly by smoothing the images with a Gaussian kernel, which is equivalent to multiplying the visibilities with a Gaussian. Therefore, by smoothing the images, we effectively down-weighted the longer baselines and reduced the noise. As shown in Fig. 6, the noise deceases rapidly with increasing smoothing scales up to few arcmin. For larger smoothing scales, the corresponding part of the uv plane is well sampled and therefore the rate of noise suppression decreases. The signal strength also decreases due to smoothing, but not as significantly as the noise (see Fig. 1). We find that the best SNR in the case of LOFAR is achieved when images are smoothed on a scale of 12 arcmin FWHM.
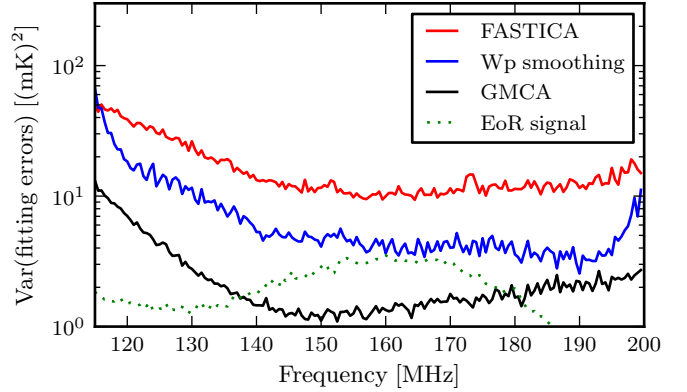

Figure 7. Variances of the foreground fitting errors (i.e. the difference between originally simulated and reconstructed foregrounds) for three removal methods for 600 h and 0.5 MHz integration, compared with the EoR signal.

3.7 Foreground removal

The success of EoR experiments depends strongly on the accuracy of the foreground removal. As the foregrounds are 2 to 3 orders of magnitude larger than the EoR signal, even small errors in their removal can severely affect the extraction of the underlying EoR signal. Foreground removal schemes are based on the assumption that the foregrounds are smooth along the frequency dimension, whereas the signal and noise are not. The signal is not expected to be smooth in frequency because it varies in space. Below we briefly describe the three foreground removal methods used in this paper.

(i) Generalized Morphological Component Analysis (GMCA): GMCA is a general source separation technique which utilizes morphological diversity and sparsity to identify different components in the data. The GMCA implementation of Chapman et al. (2013) finds a basis set in which spectrally smooth foreground components are sparsely represented and can hence be distinguished from the EoR signal and noise.

(ii) Wp smoothing: Wp smoothing (Mächler 1995) was used by Harker et al. (2009) as an EoR foreground removal algorithm. It minimizes the sum of the squared difference between the foregrounds and the data, subject to a penalty on relative changes of curvature.

(iii) FASTICA: FASTICA is an independent component analysis technique and it was implemented by Chapman et al. (2012) as a foreground removal algorithm in the context of the EoR. It separates statistically independent components of the foregrounds by maximizing non-gaussianity of their mixture.

In Fig. 7, we compare the variance of the foreground fitting errors $\text{Var}(f - \hat{f})$ for the three removal methods, where f and $\hat{f}$ are the originally simulated and reconstructed foregrounds, respectively. GMCA performs best among the three methods. We would like to note that further optimization might be possible for each of these methods. For the purpose of this paper, however, we choose GMCA to demonstrate the results.
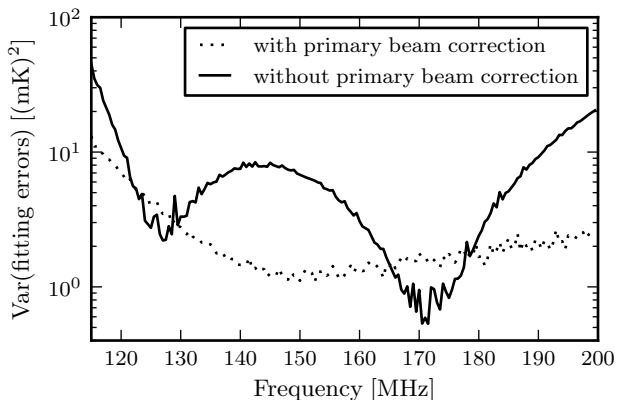
Figure 8. Comparison of variances of the GMCA foreground fitting errors with and without primary beam correction for 600 h and 0.5 MHz integration. The beam correction reconstructs the frequency coherence of the foregrounds and hence improves the foreground fitting.

We find that the primary beam correction improves the performance of the foreground removal. Fig. 8 shows that the foreground residuals are significantly reduced when the beam correction is applied. The EoR and foreground simulations were multiplied with the primary beam in the case of no correction. The GMCA was run to find two independent components (see Chapman et al. (2013) for details). Due to the frequency dependent primary beam response, GMCA fails to capture the frequency coherence of the foregrounds when the beam correction is not applied. Instead, it tries to optimize the foreground residuals in two localized parts of the bandwidth as shown in Fig. 8.

We have not considered removal of the polarized foregrounds separately, because the total intensity of the polarized foregrounds is smooth in frequency, these should be removed by the above algorithms. However, imperfect calibration of the instrumental polarization would lead to leakage of the polarized foregrounds into the total intensity. Such leakage would have frequency-dependent structure (Jelić et al. 2010) and therefore it may not be removed by the above algorithms. We assume that the instrumental polarization will be well calibrated.

4 VARIANCE MEASUREMENT AND PARAMETER ESTIMATION

Our interest lies in extracting from the data the variance of the signal as a function of frequency. The variance of the data can be measured in the image plane as $\langle (X - \bar{X})^2 \rangle$ where X is the flux density measured at a pixel, and $\bar{X}$ is the average flux density in the image. We measure the variance of the simulated data at every spectral channel in the image plane because the EoR signal and the foregrounds are simulated in the image plane. However, it is preferable to measure it in the uv plane for the actual observations because they are measured as visibilities.

The foreground extracted data d contain the EoR signal s, the noise n and foreground fitting errors r. The variance estimator can be expressed as

$$\text{Var}(d_\nu) = \langle d_\nu^2 \rangle = \langle (s_\nu + n_\nu + r_\nu)^2 \rangle, \tag{11}$$

where the subscript ν indicates the spectral channel. We have assumed in the above equation that the mean value of the data is zero, as is the case in interferometric images made with no zero spacing. The above equation can be further expanded as

$$\begin{aligned} \text{Var}(d) &= \langle s^2 + n^2 + r^2 + 2sr + 2nr \rangle \\ &= \text{Var}(s) + \text{Var}(n) + \text{Var}(r) + 2\langle rs + rn \rangle, \tag{12} \end{aligned}$$

where the first equality follows because the signal and the noise are uncorrelated so their cross-correlation $\langle sn \rangle$ is zero. The subscript ν has been dropped for convenience but all quantities are measured for each spectral channel.

The signal variance $\text{Var}(s)$ can be estimated by measuring the variance of data $\text{Var}(d)$ and subtracting from it the expected noise variance $\text{Var}(n)$, foreground fitting error variance $\text{Var}(r)$ and the cross-correlation between the noise and the foreground fitting errors $2\langle rn \rangle$. (see equation (12)). The noise and the foreground fitting errors are correlated due to the part of the noise that is removed by the foreground removal algorithm. We do not correct for the cross-correlation between the signal and the foreground fitting errors $2\langle rs \rangle$ because it will not be possible to estimate it in the case of actual observations from the data. However, we believe it would not be very significant because we find from the simulations that the term $2\langle rs \rangle$ is much smaller than other terms in equation (12). For the purpose of our simulations, variances of the noise, the foreground fitting errors and their cross-correlation are estimated from many Monte Carlo realizations of the noise and the foregrounds. In the case of actual observations, such noise realizations will be obtained from the data by differencing consecutive spectral channels of very narrow bandwidths (12 KHz). The foregrounds and the EoR signal, being smooth on these scales, are expected to get subtracted. However, since the noise is uncorrelated in different spectral channels, channel differencing is expected to yield good estimates of the noise. Estimating the foreground fitting errors from the data might be difficult in the case of actual observations and we may have to rely on foreground simulations for this purpose. However, one possible way might be to split the data into two equal integration time intervals, run the foreground removal on the two data sets separately and then subtract the residuals of one from another. The signal will get subtracted in this operation and the noise estimate can be subtracted from the variance of the remaining to obtain the estimate of the fitting errors.

The statistical error in the variance estimation can be obtained by computing the variance of the the variance estimator. For Gaussian noise, it is given by (Casella & Berger 2002)

$$Var(\hat{\sigma}^2) = \frac{2\sigma^4}{N}, \tag{13}$$

where $\hat{\sigma}^2$ is the variance estimator, σ is the true noise RMS and N is the number of measurement samples. Here, σ is the noise RMS after the primary beam correction has been applied. As the noise increases away from the centre of the image, a larger image size implies a higher σ and therefore a larger error. On the other hand, larger the image size, larger the number of independent samples N for the variance measurement. Considering these two counteracting effects,

we determine the image size for variance measurement that minimizes the error. This leads to a frequency dependent image dimension, but takes advantage of the larger field of view at lower frequencies.

4.1   Cross-variance measurement

An alternative way to measure the variance is to cross-correlate consecutive spectral channels. We will refer to such measurement 'cross-variance'. The cross-variance is measured as $\langle X_i X_{i+1} \rangle$ where X_i and X_{i+1} are flux densities at the same pixel in i^{th} and $(i+1)^{\text{th}}$ spectral channels, respectively. Measuring the cross-correlation of channels separated in frequency by $\Delta\nu$ is equivalent to measuring the power spectrum at a single scale parallel to the line of sight $k_\parallel = 2\pi/\Delta\nu$. Therefore, the cross-variance gives the variance in frequency direction on a single $k_\parallel$ mode.

The cross-variance estimator can be expressed as

$$
\begin{aligned}
\langle d_i d_{i+1} \rangle &= \langle (s_i + n_i + r_i)(s_{i+1} + n_{i+1} + r_{i+1}) \rangle \\
&= \langle s_i s_{i+1} + r_i r_{i+1} \rangle + 2\langle s_i r_{i+1} + r_i n_{i+1} \rangle, \quad (14)
\end{aligned}
$$

where we have assumed the pairs of cross terms such as $\langle s_i r_{i+1} \rangle$ and $\langle r_i s_{i+1} \rangle$ to be equal. The advantage of estimating the cross-variance as compared to the variance is that since noise is uncorrelated between different spectral channels, the cross term $\langle n_i n_{i+1} \rangle$ averages to zero. Similar to variance estimation, the signal and noise are uncorrelated, so their cross-correlation $\langle s_i n_{i+1} \rangle$ is zero. And the cross-correlation between the signal and the foreground fitting errors $\langle s_i r_{i+1} \rangle$ is not corrected for. Therefore, the signal cross-variance $\langle s_i s_{i+1} \rangle$ is estimated by measuring the cross-variance of the data $\langle d_i d_{i+1} \rangle$ and subtracting from it the foreground fitting error cross-variance $\langle r_i r_{i+1} \rangle$ and the cross-variance of the noise and the foreground fitting errors $2\langle r_i n_{i+1} \rangle$.

Unlike for variance estimation, the noise variance does not need to be subtracted by hand in the case of cross-variance estimation, reducing the chance for the systematic errors. Additionally, the statistical error in the measurement reduces by a factor of 2 in variance as we will show in equations (15) and (16). The variance of the cross-variance estimator $\hat{\sigma_c}^2$ is given by

$$
\begin{aligned}
\text{Var}(\hat{\sigma_c}^2) &= \text{Var}[\text{E}(X_i X_{i+1})] = \text{E}[\text{Var}(X_i X_{i+1})] \\
&= \text{E}\left\{ \text{E}\left[(X_i X_{i+1})^2 \right] - [\text{E}(X_i X_{i+1})]^2 \right\} \\
&= \text{E}\left\{ \text{E}\left[(X_i X_{i+1})^2 \right] \right\}, \quad (15)
\end{aligned}
$$

where the last equality follows because the noise in two different spectral channels i.e. X_i and X_{i+1} are uncorrelated. For the same reason, equation (15) can be further simplified as

$$
\text{Var}(\hat{\sigma_c}^2) = \text{E}\left[\text{E}(X_i^2)\text{E}(X_{i+1}^2) \right] = \frac{\sigma_i^2 \sigma_{i+1}^2}{N}, \quad (16)
$$

where σ_i and σ_{i+1} are the RMS noise in the ith and $(i+1)^{\text{th}}$ spectral channels respectively.

The cross-variance of the signal is slightly lower than its variance at 12 arcmin (FWHM), 0.5 MHz resolution scale. This is because small-scale structures which are coherent on scales smaller than 1 MHz do not contribute to the cross-variance measurement. However, the aforementioned advantages of cross-variance estimation supersede this disadvantage, as we will show in Section 5.

4.2   Parameter estimation

Once the signal variance has been extracted, we fit the model described in Section 2 to it and estimate the best-fitting parameters. We used Markov Chain Monte Carlo (MCMC) technique to explore our 4D parameter space - $z_r, \Delta z, \beta, A$. MCMC maps the posterior probability distribution $P(\theta|D)$ of the model parameters θ given the observed (here simulated) data D. The best-fitting parameters are obtained at the point in parameter space where the posterior is maximized. The posterior can be obtained from the likelihood $P(D|\theta)$ and the prior $P(D)$ by Bayes' theorem:

$$
P(\theta|D) \propto P(D|\theta)P(D). \quad (17)
$$

We assume uniform priors and therefore mapping the posterior is the same as sampling the likelihood. We used the code EMCEE (Foreman-Mackey et al. 2013) to map the likelihood and ultimately find the maxima. The code uses multiple random walkers to sample the likelihood function. At each step, the likelihood is computed assuming Gaussian noise as

$$
P(D|\theta) = \prod_i \frac{1}{\sqrt{2\pi\sigma_{n,i}^2}} \exp\frac{-(D_i - M_i(\theta))^2}{2\sigma_{n,i}^2}, \quad (18)
$$

where $D_i, M_i(\theta)$ are the measured and predicted (by model) variance values for the ith spectral channel, $\sigma_{n,i}^2$ is the variance of the error in measurement as given by equation (13) or (16).

In order to sample the parameter space, EMCEE iteratively draws samples for each random walker using a proposal distribution based on the current positions of other walkers. If the likelihood at the proposed position is higher than the current one, the step is accepted. If it is lower than the current likelihood, it is accepted with a certain probability. Over time, the chain explores the parameter space and maps the likelihood function. To obtain the posterior distribution of reionization parameters, we marginalize over the remaining model parameters i.e. the scaling amplitude A and the power law index β as

$$
P(z_r, \Delta z|D) = \int P(z_r, \Delta z, A, \beta|D)\, dA\, d\beta. \quad (19)
$$

In principle, we should use the chi-squared distribution for the likelihood function, because the error on the variance measurement of a Gaussian noise follows a chi-square distribution. However, we get almost the same results for both Gaussian and chi-squared distributions. This is the case because a chi-square distribution with large degrees of freedom converges to a Gaussian distribution, suggesting that the number of samples in our measurement suffices for the central limit theorem to hold.

5   RESULTS

The simulation pipeline described in Section 3 generates the mock observational data sets. We estimate the signal variance from the mock data and fit the model described in Section 2 to it. Fig. 9 shows results of the model fitting and
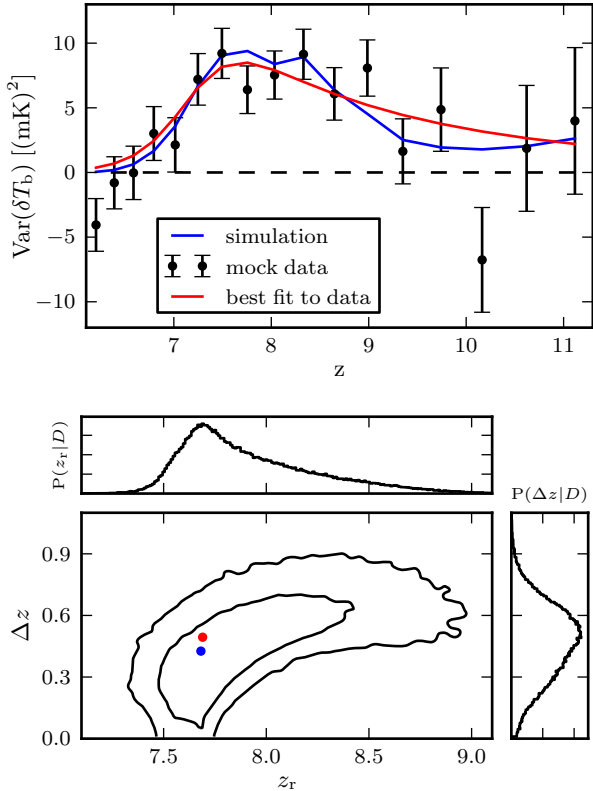
Figure 9. Model fitting and parameter estimation results for a mock data set of 600 h of integration. Top panel: Variance of $\delta T_{\rm b}$ as a function of redshift. Blue curve: original EoR signal; black points: estimated variance; red curve: model fit to the data. Error bars are calculated from equation (13). Bottom panel: constraints on the reionization parameters with 68 and 95 per cent confidence levels. Blue and red dots show the actual and best-fitting values of the parameters respectively. Marginalized probability densities are plotted on the sides.

Figure 10. Marginalized probability densities for the power-law index β and the scaling amplitude A

Figure 11. Cross-variance extraction (obtained by cross-correlating adjacent frequency channels) and parameter estimation for 600 h of integration. The details are the same as for Fig. 9 except that the error bars are calculated from equation (16). The constraints are improved as compared to the variance measurement because the noise in different channels is uncorrelated.

parameter estimation for a realization of data. The top panel shows the actual and estimated signal variance and the best model fit to the data of 600 h of integration. Foreground removal and variance measurement are performed at 0.5 MHz resolution but the variance measurements are then averaged to 5 MHz. Such re-binning of the data is done only for the convenience of showing the results but it does not affect the model fitting. The horizontal dashed line of zero variance is drawn to illustrate the significance of the detection. The bottom panel shows the obtained marginalized posterior probability densities of the reionization parameters. The actual values of the used EoR simulation parameters were $z_r = 7.68$ and $\Delta z = 0.43$. The extracted parameter values are $z_r = 7.72^{+0.37}_{-0.18}$ and $\Delta z = 0.53^{+0.12}_{-0.23}$. The errors are given at 68 per cent confidence.

Fig. 10 shows the marginalized probability density functions (PDF) for the power law index β and the scaling amplitude A. It is clear that β is not constrained well with 600 h of integration. This is due to the poor sensitivity and the limitation of the model at high redshifts. However, the uncertainty in β does not significantly affect the variance during the peak of reionization and hence it does not affect the estimation of the EoR parameters $z_{\rm r}$ and Δz.
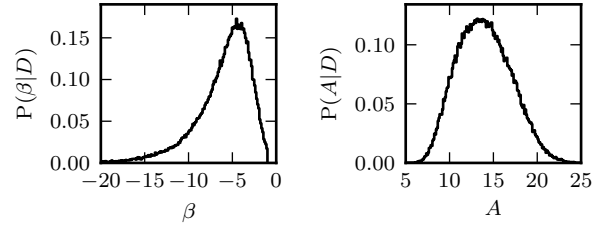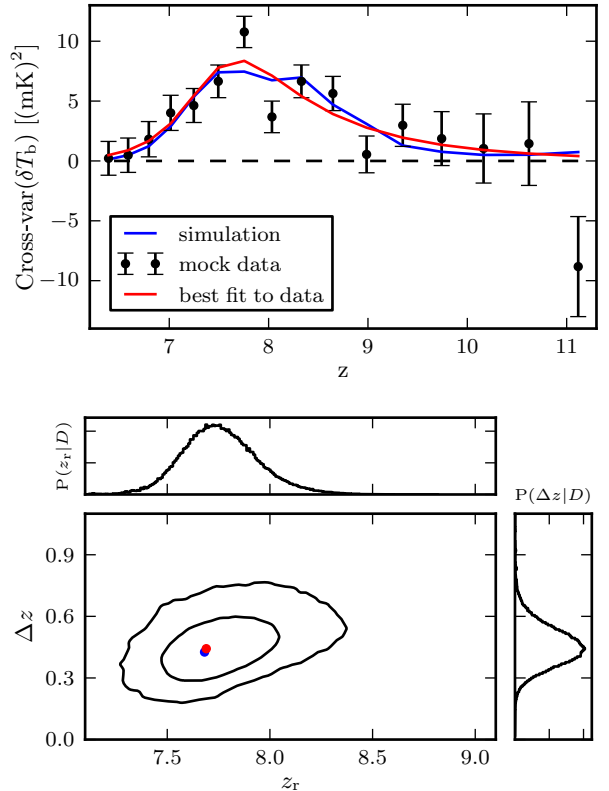
Fig. 11 shows the results of the cross-variance extraction. The constraints on the EoR parameters in the bottom panel are plotted on the same scale as in Fig. 9 to illustrate that the cross-variance measurement yields tighter constraints. The extracted values of the EoR parameters are: $z_r = 7.73^{+0.20}_{-0.16}$ and $\Delta z = 0.44^{+0.10}_{-0.09}$.

The parameter estimation results show that LOFAR can not only detect the EoR signal with 600 h of integration, but that it can also constrain the redshift and duration of reionization, assuming that the calibration errors and all the systematics are under control. Measurement of the signal variance can be used to distinguish between inside-out and outside-in models of reionization (Watkinson & Pritchard
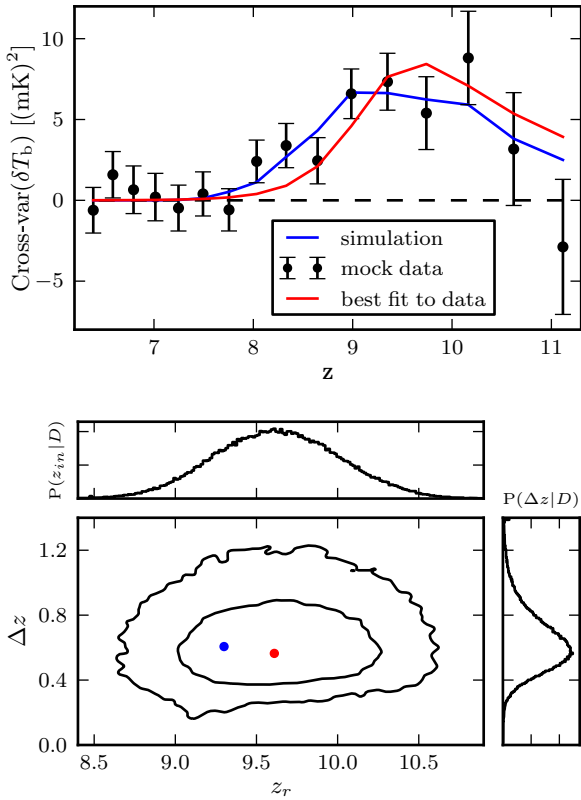
Figure 12. Cross-variance extraction for a different history of the EoR, where reionization was completed earlier. Due to the higher system noise at low frequencies, the errors on the estimated parameters become larger if reionization were to occur at higher redshifts.

2013). Measuring the redshift and duration of reionization will provide important constraints for the simulations and theoretical models, which then can improve our understanding of the EoR and the first objects in the Universe.

5.1   A different reionization history

The results shown in Fig. 9 and Fig. 11 indicate that LOFAR can in principle detect and constrain the EoR for a particular history of reionization. However, the exact redshift of reionization is unknown. Therefore, it is necessary to test whether it would be possible to constrain the EoR parameters for a different history of reionization. In particular, the signal detection may become more difficult if reionization was completed at higher redshifts because the system noise increases at lower frequencies. We therefore perform the same exercise of parameter estimation for another simulation with $z_r = 9.30$ and $\Delta z = 0.61$. The results are shown in Fig. 12. The extracted values of the parameters are: $z_r = 9.60^{+0.41}_{-0.37}$ and $\Delta z = 0.55^{+0.21}_{-0.13}$. For the same quality of the data (i.e. 600 h and 0.5 MHz integration), we obtain poorer constraints on the EoR parameters as expected due to the higher noise at lower frequencies, but it would still be possible to determine the timing and duration of the EoR.

5.2   Better quality data

The LOFAR-EoR project plans to acquire a few thousand hours of data over the coming years in order to constrain and understand the process of reionization. Therefore, it is important to check whether the proposed model for the variance statistic with its limitations would work for better quality (higher SNR) data in the future with LOFAR or the Square Kilometer Array (SKA) (Mellema et al. 2013). For this purpose, we simulated observational data sets for 1200 h of integration, keeping all other parameters same as before. Fig. 13 shows the results of the cross-variance measurement for such a data set. The extracted values of the parameters are: $z_r = 7.71^{+0.13}_{-0.11}$ and $\Delta z = 0.44^{+0.07}_{-0.09}$. As expected, the constraints on the reionization parameters have improved because of the reduced noise. However, some of the systematic errors become significant, which we otherwise neglected for the 600-hour case. The signal variance is under-estimated. This is due to the part of the signal that is removed by the foreground removal algorithm. We ignored the correlation between the signal and the foreground fitting errors, but with 1200 h of integration, the error bars become small enough for the effect of this correlation to become noticeable. The systematic errors introduced do not bias the estimates of the reionization parameters as shown in the bottom pannel of Fig. 13. However, they would affect the power spectrum estimation. Therefore, the foreground removal algorithms would need to be improved for an accurate extraction of the signal.

The statistical errors at higher redshifts, where the model does not describe the signal well, remain too large to affect the fitting. But it can be seen that with ∼2000 h of integration, the mismatch between the model and the signal at the beginning of reionization would become a source of systematic errors. We hope that the development in theory and simulations over the coming years will enable us to use improved models to describe the history of reionization.

6   CONCLUSIONS

We have investigated the extraction of the variance of the redshifted 21-cm emission as a tool to detect and constrain the global history of reionization. We have used an end-to-end simulation of the LOFAR-EoR case to demonstrate that the variance measurement is a promising tool for EoR experiments.

We parametrized the evolution of the EoR signal variance with redshift in terms of four model parameters including a characteristic redshift and duration of reionzation. We then generated realistic mock observations using an end-to-end simulation pipeline to test the variance statistics. This study has helped us realize the implications of instrumental characteristics such as uv coverage, primary beam response on the observations and to investigate different strategies for data analysis. For instance, we have quantified the impact of the time averaged frequency dependent primary beam on the foreground removal.

We have shown that LOFAR should be able to achieve a significant detection of the EoR signal in 600 h of integration on a single field, assuming that the calibration errors are small and the point sources can be removed properly.
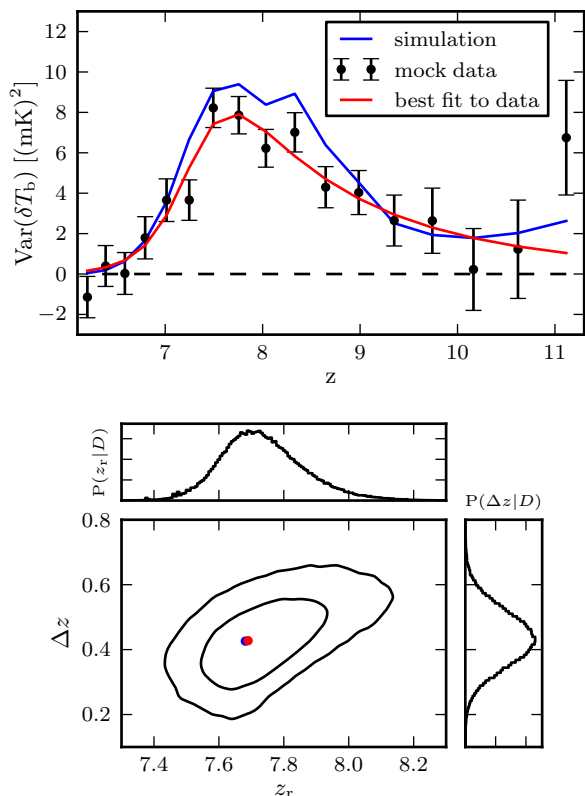
Figure 13. Variance extraction and parameter estimation results for 1200 h of integration. The constraints on reionization parameters improve with better signal-to-noise ratio, but the systematic error due to the part of the signal that is removed by foreground removal starts to become significant in the model fitting.

Additionally, it should be able to constrain the redshift z_r and duration Δz of reionization. We used a simulation with $z_r = 7.68$ and $\Delta z = 0.43$ to test our parameter estimation pipeline, and have been able to extract the parameters as $z_r = 7.72^{+0.37}_{-0.18}$ and $\Delta z = 0.53^{+0.12}_{-0.23}$, where the error bars are given at 68 per cent confidence. For the same quality of the data, we are able to get better constraints by measuring the cross-variance i.e. the cross-correlation between consecutive spectral channels. The cross-variance improves the results because the noise in different spectral channels is uncorrelated, and hence it averages to zero. For 600 h of integration, we have been able to extract the EoR parameters as $z_r = 7.73^{+0.20}_{-0.16}$ and $\Delta z = 0.44^{+0.10}_{-0.09}$ with cross-variance measurement.

If reionization was completed at higher redshifts where the sensitivity of the telescope decreases, the constraints on the EoR parameters would be poorer. For a simulation with $z_r = 9.30$ and $\Delta z = 0.61$, we have been able to constrain the reionization parameters as $z_r = 9.60^{+0.41}_{-0.37}$ and $\Delta z = 0.55^{+0.21}_{-0.13}$.

The constraints on the EoR parameters will improve with more hours of integration. For 1200 h of integration and a simulation with $z_r = 7.68$ and $\Delta z = 0.43$, we have been able to constrain the EoR parameters as $z_r = 7.71^{+0.13}_{-0.11}$ and $\Delta z = 0.44^{+0.07}_{-0.09}$. However, the systematic errors due to the part of the signal that is removed by the foreground removal algorithm will become significant with better quality

data. Therefore, the foreground removal algorithms and the models of reionization would then need improvements for an accurate extraction of the EoR signal.

We intend to explore further possibilities of optimization in the data analysis and foreground extraction. Eventually, the goal is to apply the variance statistic to LOFAR-EoR observations.

REFERENCES

Beardsley A. P. et al., 2013, MNRAS, 429, L5
Bernardi G. et al., 2009, A&A, 500, 965
—, 2010, A&A, 522, A67
Bittner J. M., Loeb A., 2011, JCAP, 4, 38
Bolton J. S., Becker G. D., Wyithe J. S. B., Haehnelt M. G., Sargent W. L. W., 2010, MNRAS, 406, 612
Bowman J. D., Morales M. F., Hewitt J. N., 2006, ApJ, 638, 20
Casella G., Berger R. L., 2002, Statistical Inference, Duxbury advanced series in statistics and decision sciences. Thomson Learning
Chapman E. et al., 2013, MNRAS, 429, 165
—, 2012, MNRAS, 423, 2518
Fan X. et al., 2006, AJ, 132, 117
—, 2003, AJ, 125, 1649
Field G. B., 1959, ApJ, 129, 536
Finkelstein S. L. et al., 2012, ApJ, 758, 93
Foreman-Mackey D., Hogg D. W., Lang D., Goodman J., 2013, PASP, 125, 306
Furlanetto S. R., Oh S. P., Briggs F. H., 2006, PhysRep, 433, 181
Harker G. et al., 2010, MNRAS, 405, 2492
—, 2009, MNRAS, 397, 1138
Hinshaw G. et al., 2013, ApJS, 208, 19
Iliev I. T., Mellema G., Pen U.-L., Bond J. R., Shapiro P. R., 2008, MNRAS, 384, 863
Iliev I. T., Mellema G., Shapiro P. R., Pen U.-L., Mao Y., Koda J., Ahn K., 2012, MNRAS, 423, 2222
Jelić V., Zaroubi S., Labropoulos P., Bernardi G., de Bruyn A. G., Koopmans L. V. E., 2010, MNRAS, 409, 1647
Jelić V. et al., 2008, MNRAS, 389, 1319
Labropoulos P. et al., 2009, ArXiv e-prints
Mächler M., 1995, The Annals of Statistics, 23, 1496
Madau P., Meiksin A., Rees M. J., 1997, ApJ, 475, 429
McQuinn M., Zahn O., Zaldarriaga M., Hernquist L., Furlanetto S. R., 2006, ApJ, 653, 815
Mellema G. et al., 2013, Experimental Astronomy, 36, 235
Mesinger A., Furlanetto S., 2007, ApJ, 669, 663
Mesinger A., Furlanetto S., Cen R., 2011, MNRAS, 411, 955
Morales M. F., Hewitt J., 2004, ApJ, 615, 7
Ono Y. et al., 2012, ApJ, 744, 83
Parsons A. R. et al., 2010, AJ, 139, 1468
Pen U.-L., Chang T.-C., Peterson J. B., Roy J., Gupta Y., Bandura K., 2008, in American Institute of Physics Conference Series, Vol. 1035, The Evolution of Galaxies Through the Neutral Hydrogen Window, Minchin R., Momjian E., eds., pp. 75–81
Pentericci L. et al., 2011, ApJ, 743, 132
Pritchard J. R., Loeb A., 2008, Phys. Rev. D, 78, 103511
Schenker M. A., Stark D. P., Ellis R. S., Robertson B. E.,

Dunlop J. S., McLure R. J., Kneib J.-P., Richard J., 2012, ApJ, 744, 179

Shaver P. A., Windhorst R. A., Madau P., de Bruyn A. G., 1999, A&A, 345, 380

Taylor G. B., Carilli C. L., Perley R. A., eds., 1999, Astronomical Society of the Pacific Conference Series, Vol. 180, Synthesis Imaging in Radio Astronomy II

Theuns T., Schaye J., Zaroubi S., Kim T.-S., Tzanavaris P., Carswell B., 2002, ApJL, 567, L103

Thomas R. M. et al., 2009, MNRAS, 393, 32

Tingay S. J. et al., 2013, PASA, 30, 7

van Haarlem M. P. et al., 2013, A&A, 556, A2

Wang F. Y., 2013, A&A, 556, A90

Watkinson C. A., Pritchard J. R., 2013, ArXiv e-prints

Zahn O. et al., 2012, ApJ, 756, 65